



The Palaeoclimate and Terrestrial Exoplanet Radiative Transfer Model Intercomparison Project (PALAEOTRIP): experimental design and protocols

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Abstract. Accurate radiative transfer calculation is fundamental to all climate modelling. For deep palaeoclimate, and increasingly terrestrial exoplanet climate science, this brings both the joy and the challenge of exotic atmospheric compositions. The challenge here is that most standard radiation codes for climate modelling have been developed for modern atmospheric conditions and may perform poorly away from these. The palaeoclimate or exoclimate modeller must either rely on these or use bespoke radiation codes, and in both cases rely on either blind faith or ad hoc testing of the code. In this paper, we describe the protocols for the Palaeoclimate and Terrestrial Exoplanet Radiative Transfer Model Intercomparison Project (PALAEOTRIP) to systematically address this. This will compare as many radiation codes used for palaeoclimate or exoplanets as possible, with the aim of identifying the ranges of far-from-modern atmospheric compositions in which the codes perform well. This paper describes the experimental protocol and invites community participation in the project through 2017–2018.

1 Introduction

Earth's atmospheric composition has varied dramatically through time, and yet-to-be-discovered terrestrial exoplanets will add untold diversity. An example model of late Archean atmospheric composition would be of 30 000 ppmv CO₂, 1000 ppmv CH₄, with no oxygen or ozone, and with an unknown nitrogen inventory, whereas escape from “snowball Earth” glaciation may take 10% CO₂. A fundamental part of the palaeoclimate problem, and equivalently the exoclimate

problem, may be stated as follows: given some atmospheric composition, what was the energy balance of the planet? Or, for given atmospheric composition and incident solar flux, what was the surface temperature?

Conceptually, to solve this, the atmospheric composition and structure, and the surface properties must be simulated, the equations for which are well known (e.g. Goody and Yung, 1989). Regrettably, implementation is far from simple. Millions of gas absorption lines from numerous gases are relevant to the climate problem. Herculean work has assembled most of these into large and oft-revised databases (e.g. Rothman et al., 2013). From these databases, absorption cross sections may be calculated as a function of temperature and pressure. Even these cross sections, calculated with standard assumptions regarding the shape of absorption lines, have some notable disagreement with observations, and smoothly varying “continuum” absorption must be added to produce realistic cross sections. Armed with cross sections, the radiative transfer equations may then be solved at the natural resolution of the lines – a so-called line-by-line calculation. Alas, these can take time of the order of minutes to hours for a single column, and hence are too slow by many orders of magnitude to be used in a climate model.

In a general circulation model (GCM), the radiative transfer for a single column must be evaluated in a fraction of a second. Consequently, simplifications must be made in the treatment of the radiative transfer, and the spectral dependence must be heavily parameterized. To optimize efficiency, these parameterizations may be made for limited ranges of atmospheric composition or column abundances of absorbing molecules. Often, these parameterizations were made a decade or more ago, with poor documentation. When an

older (and likely faster) GCM is used for palaeoclimate research, one is automatically in the situation of using a legacy radiation code.

At the other end of the modelling spectrum, there is still a cottage industry of bespoke development of fast-enough radiative transfer codes for deep palaeoclimate, planetary atmospheres, or other obscure radiative transfer problems, where all the required steps are made ad hoc. However, in some cases, the resources required to sufficiently test the code are unavailable locally.

Three broad classes of problem arise. First, whilst excellent parameterization is possible within design ranges, some parameterizations do not perform as well as a third-party user may hope. For example, intercomparison of radiation codes used for the IPCC Fourth Assessment Report (Collins et al., 2006) showed that many codes simulated the changes due to a doubling of carbon dioxide poorly. Second, the regions of parameter space of interest for palaeoclimate often fall outside the design ranges of codes, so performance may deteriorate (e.g. Goldblatt et al., 2009b). Third, errors are made in parameterizations (especially in bespoke codes) which can remain undetected through review and for some years afterwards.

The palaeoclimate or exoplanet modeller is thus in a bind. The science interest is in novel atmospheric compositions, whose radiation properties are outside the intuition of most non-specialists. It would be prudent to test any fast radiation code that one planned to use against a well-trusted line-by-line code across the parameter space of interest (e.g. Goldblatt et al., 2009b; Wolf, 2013; Yang et al., 2016); however, doing this requires both the specialist knowledge in radiative transfer, the local availability of such a model and a lot of time and energy. All of these can be hard to come by.

With the Palaeoclimate and Terrestrial Exoplanet Radiative Transfer Model Intercomparison Project (PALAEOTRIP), we hope to alleviate this problem. Our aim is to test a large number of fast radiation codes, both GCM and bespoke, against line-by-line models for a wide range of conditions applicable to palaeoclimate and terrestrial exoplanet research. Such intercomparison studies have a long history in application to modern conditions and anthropogenic global change (e.g. Ellingson et al., 1991; Fouquart et al., 1991; Collins et al., 2006; Oreopoulos et al., 2012; Pincus et al., 2015, 2016) and have contributed markedly to improvements in the fidelity of radiation codes and thus the robustness of climate models. Our hope is that by applying such a systematic intercomparison process to deep palaeoclimate and exoplanets will yield similar improvements. In this paper, we describe the experimental design and protocol¹. Up-to-date project information will be available at www.palaeotrip.org throughout the project.

¹Community input on the experimental design and protocols was gathered during the open peer-review process “discussion” phase of Geoscientific Model Development.

2 Experimental design

2.1 Philosophy

Our hope is that by assembling and analysing results from many radiative transfer codes outside of modern conditions, we will both help future investigators to make an educated choice of which radiative transfer code is applicable for a particular experiment and inform model developers of opportunities for improvement of models.

The standard method of radiative transfer intercomparison is to compare model output – especially changes in fluxes in response to changes in atmospheric composition – calculated on fixed atmospheric profiles. The use of fixed profiles is essential to isolate the fidelity of the radiative transfer codes (to be evaluated) from the myriad of other processes that determine the atmospheric profile. This methodology has a long history (e.g. Ellingson et al., 1991; Fouquart et al., 1991); see Collins et al. (2006) for an in-depth discussion of this methodology. We use instantaneous (unadjusted) radiative forcings. The most modern radiative transfer intercomparison projects for IPCC class models (Pincus et al., 2016) additionally use effective radiative forcings that account for a variety of rapid adjustments in GCMs; these are not included here. Our method here corresponds to the Pincus et al. (2016) assessment of “parameterization error”.

Three groups of experiments are included, addressing changes to clear-sky properties under both a solar and a M-star spectrum, and adding clouds under the solar spectrum. These give 14 experiments in total, each of which varies a parameter of key importance for palaeoclimate and Earth-like exoplanets. The choice of parameter space represents a range of mainstream assumptions about atmospheric composition throughout Earth’s history. We have explored all of this parameter space previously: see Goldblatt et al. (2009b) and Byrne and Goldblatt (2014a, b) for well-mixed greenhouse gases, Goldblatt and Zahnle (2011b, a) for clouds, and Goldblatt et al. (2009a) for varying atmospheric pressure. One class of model atmospheres that we exclude is H₂-dominated atmospheres (Wordsworth and Pierrehumbert, 2013); as air-broadened line shapes will likely not be appropriate for these, consequently a majority of codes may not perform well (that is, these atmospheres require rather specialist treatment, beyond the scope of this intercomparison).

Participating groups should run the experiments that their models are configured for, and omit any which are not possible (or are onerous) to run. We do not expect groups to perform model development in order to participate in this project. For example, a model which had the solar spectrum hard-coded and did not include N₂O absorption would run experiments 1, 2a–b, 3–6 and 13–16. A model without clouds would omit experiments 13–16.

If some absorbing gases are missing, experiments which do not focus on these can still be run, with notes in the metadata and in discussion with the project team. As our analy-

sis will focus on forcings (change from standard conditions), comparison to the standard conditions from that model will minimize the effect of any systematic offset from missing absorbers. For example, models without oxygen or ozone absorption could still run the experiments focusing on clouds.

If, for any reason, there is a limit to the number of experiments that a group can run then experiments 1–6 should be considered “core” and prioritized. A minimal set of experiments would be 1 and 2.

All of the required input files for the project are available at www.palaeotrip.org, and as a Supplement to this paper.

2.2 Model atmosphere

2.2.1 Atmospheric profile

For simplicity, all experiments use a global annual mean (GAM) profile. This is based on a profile derived from averaging of reanalysis data by Byrne and Goldblatt (2014a). This specific profile should be used, and none substituted for it. We refer to model levels as the boundary between model layers. Experiments 1–4, 6–8 and 10 use the GAM profile unmodified, whereas experiments 5 and 9 modify it as described for experiment 5.

Radiatively active species in the atmosphere are CO₂, CH₄, N₂O, H₂O, O₃ and O₂. All mixing ratios are in parts per volume. Standard mixing ratios are 0.21 for O₂, and vertically resolved profiles supplied with in the GAM profile for H₂O and O₃. For the remaining gases, referred to as well-mixed greenhouse gases (WMGHGs), mixing ratios are supplied in Table 1.

2.2.2 Line data

Line-by-line codes should use line data from HITRAN2012 (e.g. Rothman et al., 2013).

Bespoke, GCM and legacy radiation codes will use a variety of line data. It is acceptable to submit either the most current or standard version, or a variety of versions corresponding to different applications. The model version number or name and a brief description and/or link to the full description should be included as metadata with the model output, especially the version number or name of the code.

2.2.3 Stellar fluxes

Stellar fluxes are supplied for both the Sun and an example M-star (ADLeo) for models in which these are input directly.

As with line data, for codes which use a standard stellar flux, use this standard configuration and include whatever description possible. For such codes, where it is impractical to modify the stellar flux to an M-star, perform experiments 1–6 and 12–16 only.

All experiments should use an integrated stellar flux (solar constant) of 1360 W m⁻².

2.2.4 Clouds

Experiments with both low and high clouds are included. Calculations should be done with a single profile, with a cloud fraction of unity. Clouds may be specified in different ways in different radiation codes; the nominal descriptions here should be matched as well as possible given how clouds are specified in the particular radiation code, and appropriate description provided as metadata. We emphasize that the normal implementations of clouds in participant models should be used; single scattering properties are provided only for cases where this necessarily needs to be input. There is a range of good choices of representation of cloud microphysics in models (i.e. which are different but entirely reasonable), so variation in the radiative effects of clouds may arise from these rather than error per se. Nonetheless, it is of primary interest to us how the radiative effects of clouds do vary when every attempt has been made to specify cloud physical properties equivalently.

Vertical position: if clouds are specified in a layer, low clouds should be in the 900–925 hPa layer, and high clouds in the 250–300 hPa layer. If they are specified on levels, they should be at 912.5 and 275 hPa and can be specified with minimal vertical extent (or extent not exceeding the boundaries of the layer).

Low clouds are taken to be made of liquid water droplets. Thus cloud particles are well described as Mie spheres, so consistent specification across models should be straightforward. A standard low cloud should have a water path of $W = 40 \text{ g m}^{-2}$ and effective radius of 10 μm (Goldblatt and Zahnle, 2011b). For codes which require single scattering properties, output from a Mie code is provided in the Supplement (for simplicity, a single particle radius of $r = r_{\text{eff}}$ is used) and the Henyey–Greenstein phase function should be used.

High clouds are taken to be made of ice crystals, and are thus more complicated to describe, as there are a variety of ice habits which are all non-spherical. The normal parameter to describe the size of particles is the effective diameter, D_{eff} . A standard high cloud should have a water path of $W = 20 \text{ g m}^{-2}$ and effective diameter of 80 μm (Goldblatt and Zahnle, 2011b). For codes which require single scattering properties, these are taken from the “general habit mixture” of Baum et al. (2014, see also http://www.ssec.wisc.edu/ice_models/polarization.html) and provided in the Supplement, and the Henyey–Greenstein phase function should be used.

For codes which specify cloud thickness via an optical depth τ , this can be calculated directly from the extinction efficiency, Q : $\tau = \pi r^2 n Q$, where n is the number of cloud particles in the column; n is found directly as $n = W/m$, where for liquid droplet mass is $m = \rho(4/3)\pi r^3$, given density ρ , and for ice droplets m is supplied.

Table 1. Description of experiments.

Experiment no.	Parameter	Value or description
1	Name	Standard Conditions.
	Description	–
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	None.
	Run code	PT1.
	No. of runs	1.
2	Name	WMGHG variation.
	Description	The concentration of each WMGHG is varied in series (ranges below), with the other two held at standard conditions. The lower end of each range is selected for minimal radiative significance of that gas (see Byrne and Goldblatt, 2014b). The upper limit is an arbitrary guess at an upper bound for an Earth-like planet. Models should be run with concentrations evenly spaced in log units, with two runs per one log unit (e.g. $\{1 \times 10^{-9.0}, 1 \times 10^{-8.5}, 1 \times 10^{-8.0}, \dots\}$).
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	(a) CO ₂ from 10^{-9} to 10^{-1} , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O. (b) 400×10^{-6} CO ₂ , CH ₄ from 10^{-9} to 10^{-2} , 1×10^{-6} N ₂ O. (c) 400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , N ₂ O from 10^{-9} to 10^{-2} .
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	None.
	Run code	PT2a_x, PT2b_x, PT2c_x for CO ₂ , CH ₄ and N ₂ O, respectively.
	No. of runs	17 + 15 + 15 = 47.
3	Name	WMGHG variation, high background, anoxic.
	Description	The concentration of each WMGHG is varied in series, with the other two held at high conditions potentially representative of the Archean: $30\,000 \times 10^{-6}$ CO ₂ , 300×10^{-6} CH ₄ , 30×10^{-6} N ₂ O. Absorption by atmospheric oxygen and ozone should be turned off, with all other conditions as standard. Note there is no change to the $T - p$ profile. Otherwise, as experiment 2.
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	(a) CO ₂ from 10^{-9} to 10^{-1} , 300×10^{-6} CH ₄ , 30×10^{-6} N ₂ O. (b) $30\,000 \times 10^{-6}$ CO ₂ , CH ₄ from 10^{-9} to 10^{-2} , 30×10^{-6} N ₂ O. (c) $30\,000 \times 10^{-6}$ CO ₂ , 300×10^{-6} CH ₄ , N ₂ O from 10^{-9} to 10^{-2} .
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O.
	Clouds	None.
	Run code	PT3a_x, PT3b_x, PT3c_x for CO ₂ , CH ₄ and N ₂ O, respectively.
	No. of runs	17 + 15 + 15 = 47.
4	Name	Water vapour variation.
	Description	The water vapour mixing ratio is changed by a constant factor, with all other gases as standard conditions. The range of factors is $0.01 < x < 10$, which correspond to the differences a range of saturation vapour pressures from 230 to 330 K. Models should be run with concentrations evenly spaced in log units, with four runs per one log unit.
	Spectrum	Solar.
	Profile	GAM, altered water vapour profiles.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	None.
	Run code	PT4_x.
	No. of runs	13.

Table 1. Continued.

Experiment no.	Parameter	Value or description
5	Name	Surface pressure variation.
	Description	The surface pressure is varied between 0.1 and 10 bars. This is done by multiplying the pressure vector in the GAM profile by a factor $0.1 \leq y \leq 10$, and dividing mixing ratio vectors of minor absorbing species (CO_2 , CH_4 , N_2O and O_3) by y so that the mass of each absorber is conserved. Absorption by atmospheric oxygen and ozone should be turned off, because the mass of this absorber cannot be conserved at low pressure. Models should be run with y evenly spaced in log units, with four runs per one log unit.
	Spectrum	Solar.
	Profile	GAM with modified pressure.
	WHGHG	$400 \times 10^{-6} \text{CO}_2$, $1 \times 10^{-6} \text{CH}_4$, $1 \times 10^{-6} \text{N}_2\text{O}$.
	Absorbers	CO_2 , CH_4 , N_2O , H_2O .
	Clouds	None.
	Run code	PT5_x.
	No. of runs	9.
6	Name	No oxygen or ozone absorption.
	Description	Absorption by atmospheric oxygen and ozone should be turned off, with all other conditions as standard. Note there is no change to the $T - p$ profile.
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	$400 \times 10^{-6} \text{CO}_2$, $1 \times 10^{-6} \text{CH}_4$, $1 \times 10^{-6} \text{N}_2\text{O}$.
	Absorbers	CO_2 , CH_4 , N_2O , H_2O .
	Clouds	None.
	Run code	PT6.
	No. of runs	1.
7	Name	Standard Conditions, M-star spectrum.
	Description	As experiment 1, M-star spectrum substituted for solar spectrum.
	Spectrum	M-star.
	Profile	GAM.
	WHGHG	$400 \times 10^{-6} \text{CO}_2$, $1 \times 10^{-6} \text{CH}_4$, $1 \times 10^{-6} \text{N}_2\text{O}$.
	Absorbers	CO_2 , CH_4 , N_2O , H_2O , O_3 , O_2 .
	Clouds	None.
	Run code	PT7.
	No. of runs	1.
8	Name	WMGHG variation, M-star spectrum.
	Description	As experiment 2, M-star spectrum substituted for solar spectrum.
	Spectrum	M-star.
	Profile	GAM.
	WHGHG	(a) CO_2 from 10^{-9} to 10^{-1} , $1 \times 10^{-6} \text{CH}_4$, $1 \times 10^{-6} \text{N}_2\text{O}$. (b) $400 \times 10^{-6} \text{CO}_2$, CH_4 from 10^{-9} to 10^{-2} , $1 \times 10^{-6} \text{N}_2\text{O}$. (c) $400 \times 10^{-6} \text{CO}_2$, $1 \times 10^{-6} \text{CH}_4, \text{N}_2\text{O}$ from 10^{-9} to 10^{-2} .
	Absorbers	CO_2 , CH_4 , N_2O , H_2O , O_3 , O_2 .
	Clouds	None.
	Run code	PT8a_x, PT8b_x, PT8c_x for CO_2 , CH_4 and N_2O , respectively.
	No. of runs	$17 + 15 + 15 = 47$.
9	Name	WMGHG variation, high background, anoxic, M-star spectrum.
	Description	As experiment 3, M-star spectrum substituted for solar spectrum.
	Spectrum	M-star.
	Profile	GAM.
	WHGHG	(a) CO_2 from 10^{-9} to 10^{-1} , $300 \times 10^{-6} \text{CH}_4$, $30 \times 10^{-6} \text{N}_2\text{O}$. (b) $30\,000 \times 10^{-6} \text{CO}_2$, CH_4 from 10^{-9} to 10^{-2} , $30 \times 10^{-6} \text{N}_2\text{O}$. (c) $30\,000 \times 10^{-6} \text{CO}_2$, $300 \times 10^{-6} \text{CH}_4, \text{N}_2\text{O}$ from 10^{-9} to 10^{-2} .
	Absorbers	CO_2 , CH_4 , N_2O , H_2O .
	Clouds	None.
	Run code	PT9a_x, PT9b_x, PT9c_x for CO_2 , CH_4 and N_2O , respectively.
	No. of runs	$17 + 15 + 15 = 47$.

Table 1. Continued.

Experiment no.	Parameter	Value or description
10	Name	Water vapour variation, M-star spectrum.
	Description	As experiment 3, M-star spectrum substituted for solar spectrum.
	Spectrum	M-star
	Profile	GAM, altered water vapour profiles.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	None.
	Run code	PT10_x.
	No. of runs	13.
11	Name	Surface pressure variation, M-star spectrum.
	Description	As experiment 5, M-star spectrum substituted for solar spectrum.
	Spectrum	M-star.
	Profile	GAM with modified pressure.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O.
	Clouds	None.
	Run code	PT11_x.
	No. of runs	9.
12	Name	No oxygen or ozone absorption, M-star spectrum
	Description	As experiment 4, M-star spectrum substituted for solar spectrum.
	Spectrum	M-star.
	Profile	GAM.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O.
	Clouds	None.
	Run code	PT12.
	No. of runs	1.
13	Name	Low cloud, thickness variation.
	Description	A low-altitude water cloud is added to the standard profile (experiment 1), and the liquid water path varied between 10 and 100 g m ⁻² .
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	Water cloud, effective radius 10 μm, water path {10, 15, 25, 40, 63, 100} g m ⁻² .
	Run code	PT13_x.
	No. of runs	6.
14	Name	Low cloud, effective radius variation.
	Description	A low-altitude water cloud is added to the standard profile (experiment 1), and the effective radius varied between 5 and 25 μm.
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	400×10^{-6} CO ₂ , 1×10^{-6} CH ₄ , 1×10^{-6} N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	Water cloud, effective radius {5, 7.5, 10, 12.5, 15, 20, 25} μm, water path 40 g m ⁻² .
	Run code	PT14_x.
	No. of runs	7.

Table 1. Continued.

Experiment no.	Parameter	Value or description
15	Name	High cloud, thickness variation.
	Description	A high-altitude water cloud is added to the standard profile (experiment 1), and the ice water path varied between 10 and 100 g m ⁻² .
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	400 × 10 ⁻⁶ CO ₂ , 1 × 10 ⁻⁶ CH ₄ , 1 × 10 ⁻⁶ N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	Ice cloud, effective diameter 80 μm, water path {10, 15, 25, 40, 63, 100} g m ⁻² .
	Run code	PT15_x.
No. of runs	6.	
16	Name	High cloud, effective radius variation.
	Description	A high-altitude water cloud is added to the standard profile (experiment 1), and the effective diameter varied between 20 and 120 μm.
	Spectrum	Solar.
	Profile	GAM.
	WHGHG	400 × 10 ⁻⁶ CO ₂ , 1 × 10 ⁻⁶ CH ₄ , 1 × 10 ⁻⁶ N ₂ O.
	Absorbers	CO ₂ , CH ₄ , N ₂ O, H ₂ O, O ₃ , O ₂ .
	Clouds	Water cloud, effective diameter {20, 40, 60, 80, 100, 120} μm, water path 25 g m ⁻² .
	Run code	PT16_x.
No. of runs	6.	

2.2.5 Miscellaneous details

A solar zenith angle of 60° should be used for all experiments.

The surface should be black for thermal calculations and have a grey albedo of 0.12 in solar calculations. If a combined solar and thermal calculation is performed, the separation between solar and thermal albedos should be at 3 μm.

The surface temperature is 288.24 K in all experiments.

Note that, for most experiments, a literal interpretation of the changes to atmospheric conditions will imply some physical inconsistencies: there is no change in atmospheric pressure when CO₂ mixing ratio increases to 10⁻¹, water vapour may become super-saturated, and there is no change to the $T - p$ profile when gas concentrations change. These inconsistencies are tolerated, with the philosophy of designing simple and easy-to-compare experiments which test the fidelities of the radiation codes, which is best done on fixed profiles.

2.3 Experiments

The experiments are described in Table 1.

The run code is a unique identifier for each run, which should be used as the name of the output file for each run (e.g. `runcode.dat`). These all begin PT (for PALAEOTRIP, and to avoid starting a filename with a number), followed by the number of the experiment and the run number (x) within each experiment, counting from the lowest value of any quantity varied.

2.4 Submission of results

To facilitate comparison of many codes, each of which undoubtedly has its own output format, we ask that contributing scientists reformat output into the standard plain-text format described below. These formats are simple, and we have provided MATLAB codes which will write them automatically. These scripts, and sample output files, are available at www.palaeotrip.org and included in the Supplement for this paper.

For spectrally integrated output (dimensions are in watts per squared metre, W m⁻²) the PALAEOTRIP data format consists of a plain-text file with a 12-line header that includes the metadata in Table followed by the data header describing each column, consisting of the variables in Table . Each data column is 12 characters long. The formatting codes accept model output that corresponds to either pressure levels or layers and will automatically distinguish between these (levels are the boundary between model layers). Quantities on layers and levels will be exported to separate data files but in both cases the first column will correspond to the pressure at the level or centre of the layer in pascals. The filename convention is `runcode_levels.txt` and `runcode_layers.txt` (e.g. `PT2a_1_layers.txt`, `PT2a_1_levels.txt`).

For spectrally resolved output other than from line-by-line models, where available, a separate file should be provided for each flux, with pressure levels as rows and each spectral bin as a column. There should be a 12-line header that includes the metadata in Table and a

Table 2. Model output that will be accepted by PALAEOTRIP.

Variable	Description	Unit
Quantities on levels (bold variables are required):		
plevel	pressure on levels (layer boundaries)	Pa
Fswndir	direct solar flux down	W m^{-2}
Fswndif	diffuse solar flux down	W m^{-2}
Fswdn	total solar flux down (Fswndir+Fswndif)	W m^{-2}
Fswup	solar flux up	W m^{-2}
Fswnet	net solar flux	W m^{-2}
Flwdn	thermal flux down	W m^{-2}
Flwup	thermal flux up	W m^{-2}
Flwnet	net thermal flux (Flwdn-Flwup)	W m^{-2}
Quantities on layers (all should be included if any are)		
player	pressure at layer centre	Pa
Qsolar	solar heating rate	K day^{-1}
Qtherm	thermal heating rate	K day^{-1}

Table 3. Model metadata to be included with PALAEOTRIP submissions.

Variable	Metadata description
runcode	String with the code of run (see experiment descriptions)
modelname	String with the name (and version number) of model
username	String with your name (e.g. "Colin Goldblatt")
useremail	String with your email (e.g. "czg@uvic.ca")
usernotes	String with any notes about this run

field with the flux name, then column headers of the spectral bin edges in microns and the dimension of the column. The bin edges should be those native to the model. The fluxes in each bin should be provided in watts per squared metre (W m^{-2} ; that is, the integrated flux within that spectral bin). If layer properties are provided, they likewise should be integrated within each bin such that heating rates are in kelvins per day (K day^{-1}) for each bin. The filename convention is `runcode_variable.txt` (e.g. `PT2a_1_Fswdn.txt`).

All model output should be put into a single `.zip` file called `yourname_model.zip` and can be uploaded via the PALAEOTRIP website. Include a `readme.txt` file as necessary.

For line-by-line models, spectrally resolved output should be subsampled to 1 cm^{-1} resolution. Contact the PALAEOTRIP project team directly (info@palaeotrip.org) to discuss how to submit this, as it will likely have too large a file size for our online submission system.

3 Protocol and information for contributors

The final experimental design and protocols for the PALAEOTRIP are described in this paper. These were revised following formal review and informal discussion during the discussion phase of the publication process. If you

Table 4. Proposed PALAEOTRIP timeline.

Timeframe	Activity
January 2017	Submit description or protocol paper.
January–June 2017	Review of description or protocol paper. Community feedback on experimental design.
May–June 2017	Respond to review of protocol paper and finalize protocol.
July–August 2017	Final protocol published.
August–December 2017	Contribution of radiative transfer model runs.
January–April 2018	Follow up on participants for their contributions.
May–July 2018	Analysis of model output by PALAEOTRIP team.
July–August 2018	Write results paper, circulate to co-authors.
September 2018	Co-author comments.
October 2018	Revise and submit results paper.

intend to submit model output to the PALAEOTRIP project, we ask that you register your intention at www.palaeotrip.org or contact us directly. This will ensure that models are not run in duplicate by different groups, and that your model output is expected.

The anticipated timeline of the project is in Table 4. Sadly, few deadlines survive contact with academics, but we hope that this schedule is realistic and it is our intention to keep to it. We will post any updates to www.palaeotrip.org and communicate schedule changes directly to all participating scientists.

We intend that everyone submitting unique model results will be offered authorship on the final paper. Lead authorship will be by one of the project team, who will additionally determine the order of authorship (likely project team followed by contributing scientists, listed alphabetically). This paper will be circulated amongst all co-authors prior to submission.

A motivation of this project is to find out how a variety of radiation codes perform across a range of conditions applicable to palaeoclimate and exoplanets, so that future model users may know the range of conditions across which each model is likely to be accurate. Therefore, it is essential that models are able to be identified in the final paper. The analysis will be restricted to the range of conditions specified here, as an indicator of performance in palaeo- and exoclimate studies. We have no interest in, or intention of, commenting on the fitness of any model for any other purpose. It is the responsibility of scientist submitting model results to assert that the model can be identified in the final paper.

4 Summary and discussion

PALAEOTRIP will run 14 controlled experiments addressing the radiative transfer through a subset of conditions expected through Earth's past climate, and applicable to Earth-like exoplanets. We invite community participation in the experiment. Over the course of the next year, the model runs will be performed and compared. The anticipated outcome is that the community will be better informed about the perfor-

mance of available radiative transfer codes for palaeo- and exoclimate research.

The range of conditions which we have specified experiments for is somewhat “vanilla”. It likely does not represent the full range of conditions seen in Earth’s past, and will be a tiny fraction of the parameter space for Earth-like exoplanets. This is intended to get wide participation, that is, to specify conditions which most models which derive from Earth atmospheric sciences should be capable of being run for. We anticipate that, if this intercomparison is successful, we may be able to lead a more wide-ranging intercomparison in the future.

Code and data availability. A zip file containing the GAM profile, scripts to be used to write model output into the specified format and sample output is available in a Supplement to this article. The version corresponds to this paper. Updated versions will be made available through the project website, www.palaeotrip.org, as necessary. Final model output will be available from www.palaeotrip.org and as a Supplement to the paper which will describe the results of the intercomparison.

The Supplement related to this article is available online at <https://doi.org/10.5194/gmd-10-3931-2017-supplement>.

Author contributions. CG has designed the experiment and is responsible for the scientific content herein. LK and MD have helped prepare materials and provided technical assistance.

Competing interests. The authors declare that they have no conflict of interest.

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